



UNIVERSIDAD DE  
COSTA RICA

SEP Sistema de  
Estudios de Posgrado

EFIS Escuela de  
Física

# Geodesic visualization for the Frutos metric

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(Dir. thesis: Dr. rer. nat. Francisco Frutos)

Escuela de Física

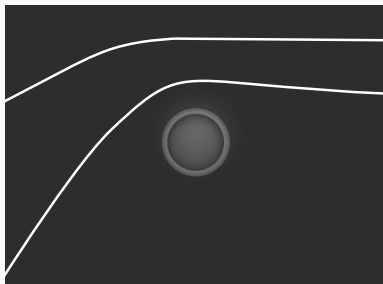
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3. Program
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# Introduction

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## Outline of the problem



- Compact/massive objects deform spacetime
- Light is affected; we are interested on the trajectories of photons
- Compare effects for different metrics, specially the Frutos metric

# Theoretical framework

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- Metrics describe the curvature of spacetime
- Geodesic equations: set of second order differential equations that contain the derivatives of the metric
- Null geodesics: trajectories of light
- Null geodesics can be used to study:
  - The properties of spacetime described by the metric<sup>1</sup>
  - High resolution gravitational lenses<sup>2</sup>
  - Radiation emitted near a compact object <sup>3</sup>

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<sup>1</sup>Hagihara (1930)

<sup>2</sup>Vincent et al. (2011)

<sup>3</sup>Yang et al. (2014), Cunningham (1975)

Objects capable of producing such deformations to spacetime

- Black holes
- Neutron stars
- Galaxies
- Dark matter

$$\frac{d^2 x^\kappa}{d\lambda^2} = -\Gamma_{\mu\nu}^\kappa \frac{dx^\mu}{d\lambda} \frac{dx^\nu}{d\lambda}$$

- $\Gamma_{\mu\nu}^\kappa$  are the Christoffel symbols
- $\lambda$  is an affine parameter
- For null geodesics (light),  $ds^2 = 0$



- Approximate metric with mass, small quadrupole moment (deformation) and rotation
- Kerr has problems when trying to match to an interior solution (e.g. Frutos 2015b)

## General structure of the metric

$$ds^2 = g_{tt}dt^2 + 2g_{t\phi}dtd\phi + g_{rr}dr^2 + g_{\theta\theta}d\theta^2 + g_{\phi\phi}d\phi^2$$

$$g_{tt} = \frac{e^{-2\psi}}{\rho^2} [a^2 \sin^2 \theta - \Delta]$$

$$g_{t\phi} = -\frac{2Jr}{\rho^2} \sin^2 \theta$$

$$g_{rr} = \rho^2 \frac{e^{2\chi}}{\Delta}$$

$$g_{\theta\theta} = \rho^2 e^{2\chi}$$

$$g_{\phi\phi} = \frac{e^{2\psi}}{\rho^2} [(r^2 + a^2)^2 - a^2 \Delta \sin^2 \theta] \sin^2 \theta$$

where  $\Delta = r^2 - 2Mr + a^2$ ;  $\rho^2 = r^2 + a^2 \cos^2 \theta$  and if  $\psi = \chi = 0$  we get the Kerr metric.

## Second order expansions for $\psi$ and $\chi$ <sup>(4)</sup>

$$\psi = \frac{q}{r^3} P_2 + 3 \frac{Mq}{r^4} P_2$$

$$\chi = \frac{qP_2}{r^3} + \frac{Mq}{r^4} \left( -\frac{1}{3} + \frac{5}{3}P_2 + \frac{5}{3}P_2^2 \right) + \frac{q^2}{r^6} \left( \frac{2}{9} - \frac{2}{3}P_2 - \frac{7}{3}P_2^2 + \frac{25}{9}P_2^3 \right)$$

where  $P_2 = (3 \cos^2 \theta - 1)/2$  (Legendre polynomial of second order).

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<sup>4</sup>Frutos (2015a)

If

$$\psi = \chi = \frac{2qM^3}{15} \frac{P_2}{r^3}$$

we obtain the metric described in Frutos et al. (2015b), already tested with our program

The parameters of these metrics are

- $M$ , mass
- $a$ , angular momentum parameter
- $q$ , quadrupole moment parameter (small)
- Geometrical units ( $G = c = 1$ )
- They can be matched to an interior solution<sup>5</sup>

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<sup>5</sup>Frutos (2015a)

- Analytical study of geodesics; properties of spacetime: Hagihara (1930)
- First computational works: Cunningham (1975), Luminet (1979)
- Semi-analytical integration: Dexter & Agol (2010), Yang (2014)
- Arbitrary numerical metric (3+1 formalism): Vincent et al. (2011)
- HPC / GPU for particular metrics: Müller (2011)
- Use of geodesics for lens reconstruction: Tessore et al. (2015)

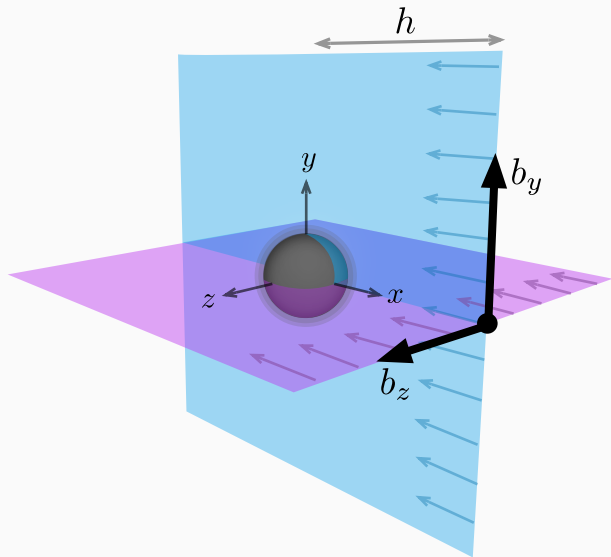
# Program

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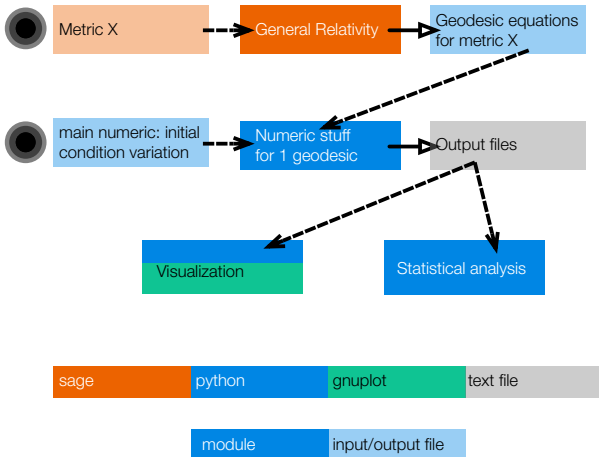
- Takes an analytical metric and computes and solves the geodesic equations
- It *should* work for *any* metric
- Python, Sage, Gnuplot
- So far applied to
  - Schwarzschild
  - Bonnor
  - Kerr
  - Frutos



# Initial conditions



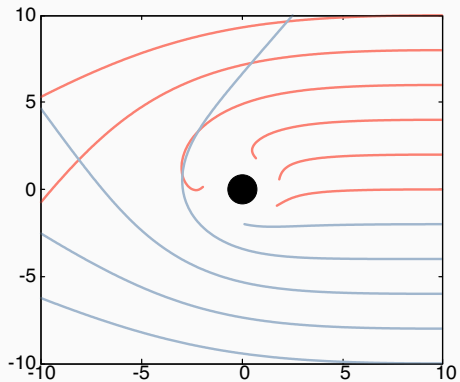
# General scheme



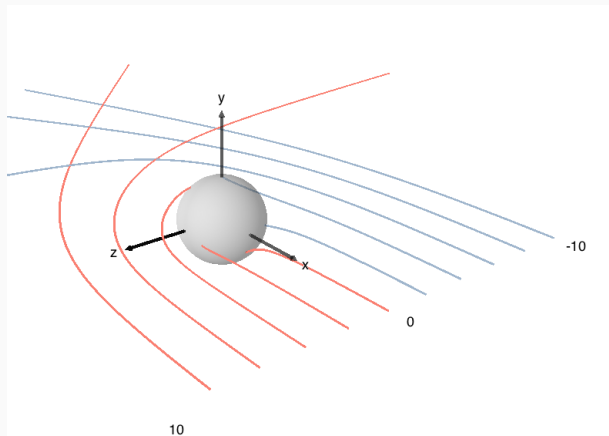
## Partial results

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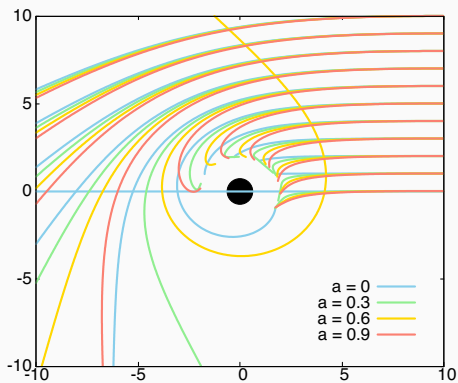
## Frutos second order: xy-plane



## Frutos second order: xz-plane

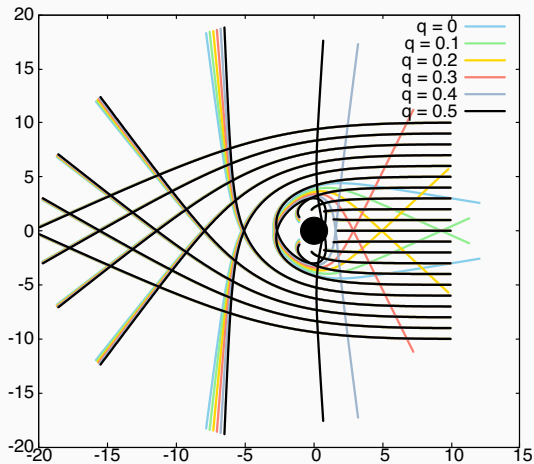


## Frutos second order: effect of the angular momentum

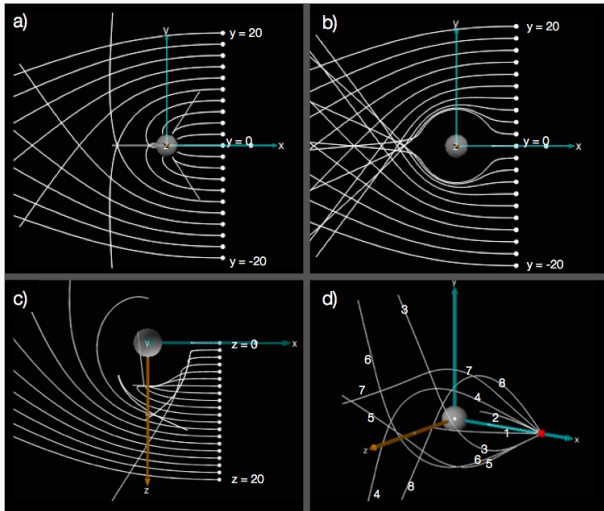


$xy$  plane,  $M = 1$ ,  $q = 0.1$

## Frutos second order: effect of the quadrupole moment



$xz$  plane,  $M = 1$ ,  $a = 0.9$











## **Future work**





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- Extend to massive particles
- Use geodesics to study emission near a compact object for different metrics (output an observable)
- Make comparisons of different metrics; effects of terms
- Writing testing modules, open-source and publish the code in a repository

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-  Müller, T. (2011). GPU-based four-dimensional general-relativistic ray tracing. *Computer Physics Communications* 183(2012) 2282-2290.
-  Vincent, F., Paumard, T., Gourgoulhon & E., Perrin, G. (2011). GYOTO: a new general relativistic ray-tracing code. *Classical and Quantum Gravity* 28, 225011.

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-  Cunningham, C. (1975). The effects of redshifts and focusing on the spectrum of an accretion disk around a Kerr black hole. *The Astrophysical Journal* 202, 788-812.
-  Luminet, J. (1979). Image of a spherical black hole with thin accretion disk. *Astronomy and Astrophysics* 75, 228-235.
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